

Distributions of H₂O and CO₂ ices on Ariel, Umbriel, Titania, and Oberon from IRTF/SpeX observations

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Abstract

We present 0.8–2.4 μm spectral observations of uranian satellites, obtained at IRTF/SpeX on 17 nights during 2001–2005. The spectra reveal for the first time the presence of CO₂ ice on the surfaces of Umbriel and Titania, by means of 3 narrow absorption bands near 2 μm . Several additional, weaker CO₂ ice absorptions have also been detected. No CO₂ absorption is seen in Oberon spectra, and the strengths of the CO₂ ice bands decline with planetocentric distance from Ariel through Titania. We use the CO₂ absorptions to map the longitudinal distribution of CO₂ ice on Ariel, Umbriel, and Titania, showing that it is most abundant on their trailing hemispheres. We also examine H₂O ice absorptions in the spectra, finding deeper H₂O bands on the leading hemispheres of Ariel, Umbriel, and Titania, but the opposite pattern on Oberon. Potential mechanisms to produce the observed longitudinal and planetocentric distributions of the two ices are considered.

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1. Introduction

The discovery of CO₂ ice on the surface of the uranian satellite Ariel (Grundy et al., 2003) raised a number of questions regarding its origin and ultimate fate. Is it a relic of Ariel's primordial volatile inventory? Was it delivered by more recent impactors? Or is it being produced in situ in chemical reactions enabled by energetic radiation? How long does exposed CO₂ ice survive on the satellite's surface, and what processes act on it? How stable is it against sublimation and sputtering? Does it participate in radiolytic chemistry, such as envisioned by Delitsky and Lane (1997, 1998), Hudson and Moore (2001), and Johnson (2001)?

Questions such as these could potentially be addressed by means of a survey of CO₂ ice as a function of longitude and planetocentric distance among the major uranian satellites, the

basic physical properties of which are listed in Table 1. The results of such a survey of the four largest satellites are presented in this paper.

2. Observations and reduction

We obtained infrared spectroscopy of uranian satellites on seventeen nights during 2001–2005 at NASA's Infrared Telescope Facility (IRTF) on Mauna Kea, as tabulated in Table 2. The data were collected using the short cross-dispersed mode of the SpeX spectrograph (Rayner et al., 1998, 2003), which covers the 0.8–2.4 μm wavelength range in five spectral orders, recorded simultaneously on a 1024 \times 1024 InSb array. Uranian satellite observations were done at an average airmass of 1.26. Maximum airmasses were 1.61, 1.77, 1.57, and 1.46 for Ariel, Umbriel, Titania, and Oberon, respectively. When observing at higher airmasses, we used an image rotator to keep SpeX's slit oriented near the parallactic angle on the sky-plane, to minimize wavelength-dependent slit losses from differential atmospheric refraction. During 2001–2002 we used the 0.5 arcsec slit, and during 2003–2005 we used the 0.3 arcsec slit, resulting in mea-

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¹ Visiting/remote observer at the Infrared Telescope Facility, which is operated by the University of Hawaii under contract from NASA.

Table 1
Major satellites of Uranus

Satellite	Distance from Uranus (10 ⁵ km)	Orbital period (days)	Radius ^a (km)	Mass ^b (10 ²³ g)	Surface gravity (cm s ⁻²)	Visual full disk albedo ^c
Miranda	1.30	1.41	236	0.66	7.9	0.307
Ariel	1.91	2.52	579	13.5	26.9	0.350
Umbriel	2.66	4.14	585	11.7	22.9	0.189
Titania	4.36	8.71	789	35.3	37.8	0.232
Oberon	5.81	13.46	761	30.1	34.7	0.205

^a Radii are from Thomas (1988).

^b Masses are from Jacobson et al. (1992).

^c Full disk albedos at 0.55 μm and 1° phase are from Karkoschka (1997), Table II.

sured spectral resolutions ($\lambda/\Delta\lambda$ for unresolved OH airglow emission lines) between 1300 and 1400 and between 1600 and 1700, respectively.

All spectra were collected as pairs, with the telescope being offset along the slit between two locations, referred to as A and B beams. Pairs of A and B spectral frames were subtracted prior to extraction using the Horne (1986) optimal extraction algorithm as implemented by M.W. Buie et al. at Lowell Observatory (e.g., Buie and Grundy, 2000).

Uranian satellite observations were interspersed between observations of nearby solar analog reference stars HD 210377 and HD 214572. We also observed more distant, but better-known solar analogs 16 Cyg B, BS 5968, BS 6060, HD 219018, and SA 112-1333. Those observations showed that HD 214572 is a reasonable solar analog in this spectral range, and that HD 210377 was also usable as a solar analog after correcting for its effective temperature, which is a few hundred K cooler than the Sun. From the solar analog observations obtained each night we computed telluric extinction and corrected the star and satellite observations to a common airmass. We then divided the airmass-corrected satellite spectra by the airmass-corrected solar analog spectra. This operation eliminated most instrumental, stellar, and telluric spectral features, and produced spectra proportional to the satellite disk-integrated albedos. Cancellation of telluric features was sometimes imperfect near 1.4 and 1.9 μm , where strong and narrow telluric H₂O vapor absorptions make sky transparency especially variable in time. Additional spurious features originate in scattered light from Uranus, such as at 0.94 μm . These could be either positive or negative, depending whether Uranus was closer to the object or sky beam position, and only affected Umbriel and Ariel spectra for longitudes near 0° or 180°. Examples of final, normalized albedo spectra are shown in Fig. 1.

Wavelength calibration was derived from telluric airglow emission lines, primarily those of OH, extracted from the satellite frames. The dispersion ranges from 0.00023 to 0.00054 μm per pixel from the shortest to longest wavelengths. The wavelength uncertainty is no more than a pixel in regions having abundant sky emission lines, which include the wavelengths where the satellites exhibit absorption features analyzed in this paper.

Our spectra could not be photometrically calibrated, since variable slit losses from different tracking time scales combine

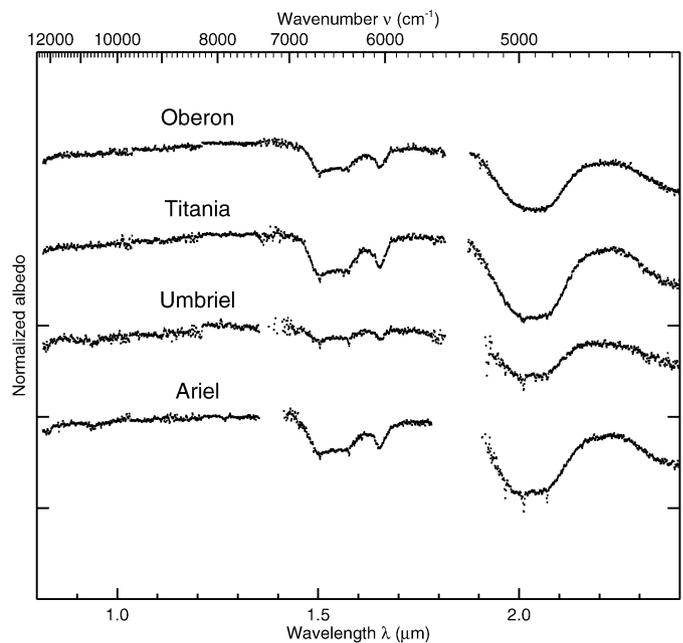


Fig. 1. Example spectra of uranian satellites, normalized at *J* band and offset upward with zero levels indicated by tick marks on the ordinate axis. Gaps around 1.4 and 1.85 μm coincide with high atmospheric opacity. The Ariel spectrum is an average of all spectra with subsolar longitudes between 210° and 330°. The Umbriel spectrum is from 2005/09/18 UT and the Titania and Oberon spectra are both from 2005/10/13 UT.

with variable focus and seeing to undermine photometric fidelity of comparisons between the faint satellites and the much brighter solar analog stars. However, other sources of photometry exist for the uranian satellites in our wavelength range (e.g., Karkoschka, 1997) and these could be used to scale the spectra to an albedo scale, enabling the use of quantitative radiative transfer models.

3. Analysis

The spectra of all four satellites in Fig. 1 look broadly similar to one another. A gentle reddish spectral slope extends from 0.8 to about 1.4 μm , followed by a water ice absorption complex between 1.45 and 1.7 μm , and two more H₂O ice absorptions between 1.9 and 2.2 μm and from 2.25 μm to the limit of our spectral coverage at 2.4 μm . The H₂O ice bands are weakest in the spectrum of Umbriel, the satellite with the lowest albedo of the four (Karkoschka, 1997). Lower albedos and weaker H₂O bands typically go together, since the same dark materials (probably carbonaceous, in the case of uranian satellites) which suppress albedos are also effective at masking H₂O ice absorptions. Although this dark material has a profound effect on the satellites' albedos, H₂O ice could still be the most abundant surface material on all four satellites, considering that thermal segregation of H₂O ice, as seen on the jovian satellites, is improbable at the low surface temperatures of the uranian satellites (Spencer, 1987). Impact gardening probably leads to intimate mixing between dark and icy components, enabling even small quantities of dark material to significantly suppress surface albedos (Clark, 1981).

Table 2
Circumstances of observations

UT date of observation mid-time	Sky conditions and H band image size	Sub-observer		Phase angle (°)	Total integration (min)
		Lon. (°)	Lat. (°)		
Ariel					
2001/07/05.59	Cirrus, 1.1''	233.8	−23.1	1.91	50
2001/07/08.61	Clear, 1.1''	304.8	−23.2	1.79	48
2002/07/16.55	Partly cloudy, 0.7''	294.8	−19.3	1.65	140
2002/07/17.56	Partly cloudy, 0.8''	79.8	−19.4	1.60	108
2003/08/05.50	Clear, 0.6''	200.0	−15.9	0.95	84
2003/08/09.51	Clear, 0.6''	53.6	−16.0	0.75	156
2003/09/07.40	Clear, 0.6''	219.8	−17.2	0.70	90
2003/10/04.24	Clear, 0.6''	93.5	−18.1	1.88	108
2003/10/08.33	Clear, 0.9''	316.6	−18.2	2.03	132
2004/07/15.50	Clear, 0.6''	159.9	−11.1	1.99	112
Umbriel					
2001/07/07.59	Cirrus, 1.0''	219.8	−23.0	1.83	52
2004/07/05.59	Thin cirrus, 0.8''	216.2	−10.9	2.32	80
2004/07/16.60	Thin cirrus, 0.8''	92.1	−11.1	1.95	74
2004/07/27.48	Clear, 0.6''	317.6	−11.4	1.51	184
2005/09/18.39	Clear, 0.6''	261.3	−9.4	0.85	196
Titania					
2001/07/06.56	Cirrus, 0.9''	237.0	−23.0	1.87	56
2001/07/07.55	Cirrus, 1.0''	277.8	−23.0	1.83	36
2003/10/08.24	Clear, 0.9''	98.0	−18.1	2.03	64
2004/07/15.60	Clear, 0.6''	213.9	−11.1	1.98	64
2005/10/13.38	Clear, 0.6''	299.6	−10.2	1.93	106
Oberon					
2001/07/06.60	Cirrus, 0.9''	164.0	−23.0	1.87	32
2001/07/08.55	Clear, 1.1''	216.2	−23.1	1.79	48
2005/10/13.27	Clear, 0.6''	110.6	−10.2	1.93	124

3.1. Water ice

The shapes of the H₂O ice absorption bands and their depths relative to one another, especially between 1.50 and 1.57 μm and around 1.65 μm, confirm that most of the H₂O ice on all four satellites is crystalline at the ~mm depths sampled by these wavelengths (e.g., Grundy et al., 1999; Hansen and McCord, 2004). Amorphous ice lacks 1.57 and 1.65 μm absorption bands and exhibits somewhat different shapes for the 1.5 and 2 μm bands (Grundy and Schmitt, 1998; Schmitt et al., 1998). The H₂O features in the satellite spectra are consistent with simple Hapke models (e.g., Hapke, 1993) having no amorphous ice at all, although we can include up to 10–20% amorphous ice before the discrepancies become conspicuous.

This predominance of crystalline ice is not surprising, since it is the more thermodynamically stable, lower energy phase. Crystalline H₂O has also been reported on other outer Solar System surfaces at comparable and greater heliocentric distances, both icy satellites and trans-neptunian objects. In fact, everywhere that the phase of H₂O ice on these objects' surfaces has been identified, it is crystalline (e.g., Buie and Grundy, 2000; Bauer et al., 2002; Grundy and Young, 2004; Jewitt and Luu, 2004).

Under certain circumstances, energetic charged particles or UV photons can disrupt the crystal structure of ice (e.g., Kouchi

and Kuroda, 1990; Johnson, 1990; Moore and Hudson, 1992; Johnson and Quickenden, 1997), but evidently this effect is so inefficient that ice on the uranian satellites is not amorphous at mm depths. Nature does tend to find its way to thermodynamically favored states such as crystallinity, in the case of H₂O ice. Thermal recrystallization goes as $e^{-E_A/kT}$, where E_A is an activation energy, k is the Boltzmann constant, and T is the temperature. Laboratory studies of crystallization rates of pure H₂O ice at much warmer temperatures (Delitsky and Lane, 1998; Jenniskens, 1998), if extrapolated over many orders of magnitude, suggest that thermally activated crystallization could be extremely slow at uranian satellite surface temperatures. However, even if the low thermal crystallization rates implied by that extrapolation are correct, the presence of impurities in the ice, or of non-thermal energy inputs from less energetic charged particles, photons, or even micrometeorites may help overcome thermal limits (e.g., Brown et al., 1978). Sputtering might also remove H₂O more efficiently than it is amorphized. Consequently, we do not subscribe to the idea proposed by Bauer et al. (2002) and by Jewitt and Luu (2004) that the existence of crystalline ice on the surfaces of these bodies implies recent heating episodes.

The H₂O ice absorption bands of jovian and saturnian satellites tend to be deeper on their leading hemispheres (e.g., Clark et al., 1984; Grundy et al., 1999; Cruikshank et al., 2005). Sim-

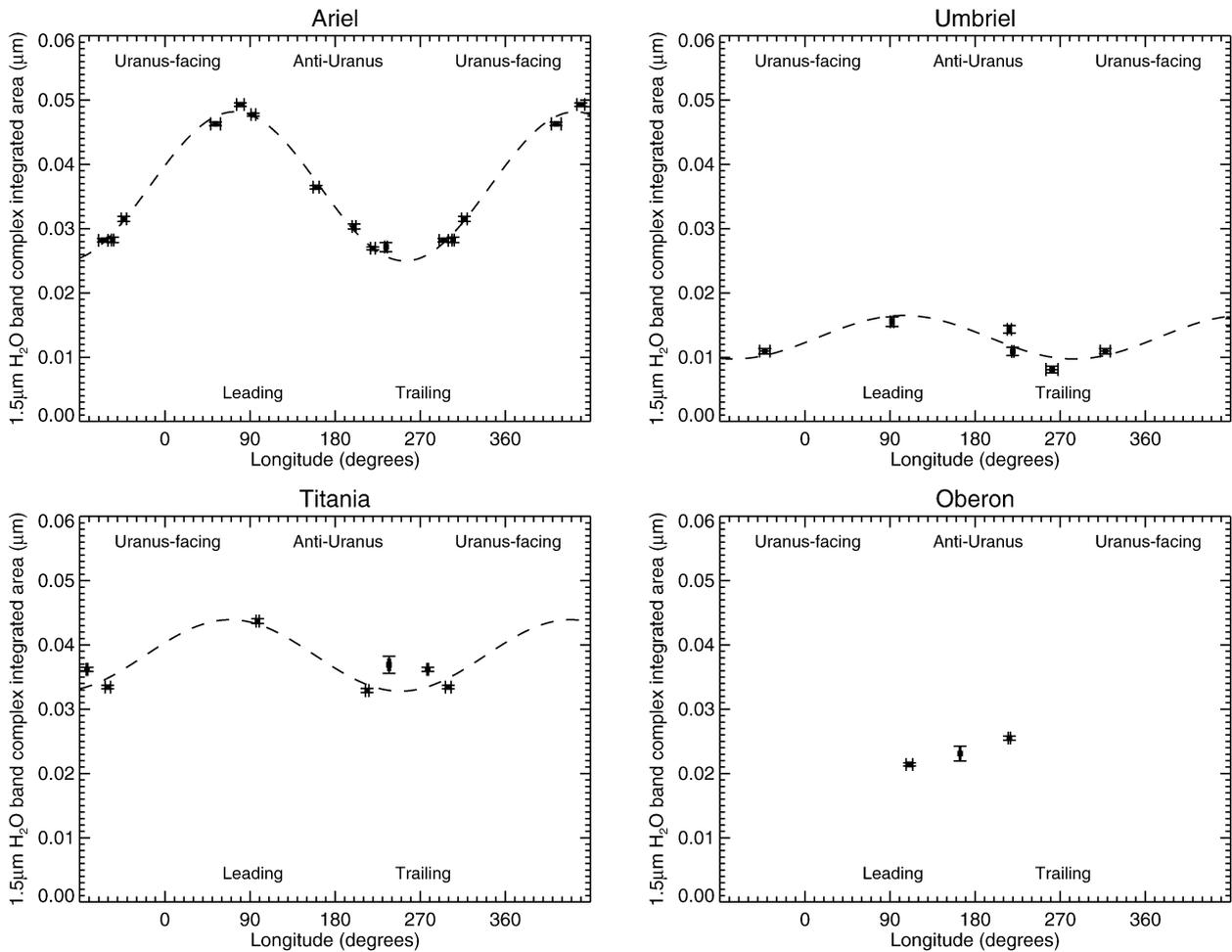


Fig. 2. Variation of integrated area of the 1.5 μm H₂O ice band complex as a function of sub-solar/sub-viewer longitude, showing large cyclical variations with more H₂O ice absorption on the leading hemispheres, at least for satellites with better longitude coverage, where sine fits to the data are plotted as dashed curves. Data are duplicated outside the 0° to 360° interval to better show the periodic trends.

ilar behavior is exhibited by the uranian satellites, except for Oberon, as shown in Fig. 2. Integrated areas of the 1.5 μm water ice absorption complex were computed by normalizing each spectrum to a line fitted to continuum wavelengths on either side of the absorption band (1.29–1.32 and 1.72–1.75 μm in this example), then integrating one minus the normalized spectrum over the band interval (1.42–1.72 μm). Ariel, with the most complete longitudinal sampling, shows a sinusoidal pattern of variation of its integrated H₂O ice bands, with maximum H₂O absorption coinciding with the leading hemisphere (90° longitude in the left-handed IAU system). Fewer longitudes were observed for the other satellites, but Umbriel and Titania both show more H₂O absorption on their leading hemispheres. The trend for Oberon appears to be reversed, but confirming observations are needed.

Magnetospheric charged particle bombardment can drive sputtering as well as radiolytic destruction of H₂O ice, and could be responsible for the observed leading-trailing asymmetries. The magnetic field of Uranus rotates rigidly, at the same frequency as Uranus spins on its axis, with a period of 17.9 h. The satellites orbit more slowly (see Table 1), so the magnetosphere overtakes them from behind and its charged particles

preferentially strike their trailing hemispheres (Cheng et al., 1986; Lanzerotti et al., 1987). Magnetospheric charged particles spiral along the magnetic field lines at gyroradii inversely proportional to the strength of the magnetic field. The dipole component of the field diminishes as the inverse cube of the distance, so gyroradii grow as the cube of the distance from Uranus. For protons, the dominant ions in the uranian magnetosphere (Bridge et al., 1986), the energies at which gyroradii match the radii of the satellites range from ~100 keV for Miranda down to <1 keV for Titania, spanning the ‘dominant’ ~10 keV energy range for magnetospheric protons (Johnson, 1990). That these protons have gyroradii larger than the radii of the satellites in the outer parts of the system suggests that outer satellites should receive more isotropically distributed proton bombardment. This tendency is moderated by the fact that the net plasma flow past the satellites, arising from the greater speed of the uranian magnetic field relative to the Keplerian motion of the satellites, is faster further from Uranus. This net plasma flow increases from 6 up to 39 km s⁻¹ from Miranda to Titania. However, these speeds correspond to proton energies only 0.2–8 eV, far below 10 keV energies. Although 10 keV protons circle field lines in orbits larger than

the radii of satellites from Umbriel on out, the orbits are executed relatively quickly compared with the rate at which the magnetic field sweeps past the satellites. Orbital periods for protons are inversely proportional to magnetic field strength, being 0.3, 1, 3, and 13 s, from Miranda through Titania. Combining these gyroperiods with net plasma flow rates, each successive proton orbit advances by 2, 14, 60, and 500 km, relative to satellites Miranda through Titania. Only at Titania and beyond does a 10 keV proton advance far enough between successive loops to begin to have much chance of looping around and hitting the satellite's leading hemisphere. Computing the actual three-dimensional spatial distribution of impacts on a satellite's surface as a function of proton energy is beyond the scope of this paper. It depends on the substantial tilt and offset of the dipole field of Uranus as well as the distribution of particle pitch angles, which could vary seasonally. However, the simple arguments presented here suggest that particles in the 10 keV range will mostly hit trailing hemispheres of satellites, and can only influence the leading hemispheres of the more distant ones. Additionally, magnetospheric plasma densities decline rapidly with planetocentric distance, more rapidly than is compensated by the increase in net plasma flow relative to satellite orbital motion (or at least they did during the Voyager 2 encounter in 1986, near the time of southern summer solstice, e.g., [Bridge et al., 1986](#); [Cheng et al., 1991](#)). Overall magnetospheric plasma sputtering rates should thus be lower on the more distant satellites. For both of these reasons, we would expect magnetospheric sputtering to produce diminishing leading-trailing contrasts with planetocentric distance. This expectation is consistent with results from sinusoidal fits to the observed longitudinal variation in H₂O bands, which give fractional variations $93 \pm 4\%$, $69 \pm 47\%$, and $34 \pm 13\%$ for Ariel, Umbriel, and Titania, respectively.

Another mechanism capable of producing leading-trailing asymmetries in the satellites' H₂O ice bands is magnetospheric delivery of carbonaceous dust particles. Dark particles from the uranian rings which become electrically charged will experience a Lorentz force from the magnetic field lines sweeping past them. This force will tend to accelerate them such that they spiral slowly outward, eventually culminating in low speed collisions onto the trailing hemispheres of the satellites, primarily the inner ones, so this mechanism could also contribute to the observed decline of leading-trailing spectral contrasts with planetocentric distance. The addition of these dark particles to the satellites' trailing hemispheres would suppress their albedos and their ice absorption bands, by reducing the fraction of photons able to escape from the surface at all wavelengths, as opposed to just the wavelengths at which ice absorbs. A possible way to establish limits on the supply of carbonaceous grains via this mechanism could involve comparison of visible wavelength albedos between leading and trailing hemispheres. Voyager images do not reveal dramatic global leading-trailing albedo patterns, but a detailed comparison, controlling for geological age by selecting specific regions, would be needed to properly address this question. Such a study is beyond the scope of this paper.

Yet another mechanism which could produce leading-trailing asymmetries diminishing with planetocentric distance is impact cratering. In general, leading hemispheres suffer more collisions with debris from outside of the Uranus system ([Zahnle et al., 2001, 2003](#)), and gravitational focusing by Uranus causes the inner satellites to be cratered more frequently and with higher average impact velocities. Large impacts could dredge up cleaner ice from below the surfaces of the satellites, resulting in deeper H₂O ice bands on leading hemispheres. Dust-sized impactors could sputter away ice and deliver exogenous, non-icy materials, creating the opposite spectral effect. Relatively little is known about the size distribution, composition, and net effect of smaller impactors on icy surfaces in the Uranus system, but the dust counter on NASA's New Horizons² spacecraft should provide useful information about the dust population in the ecliptic plane at large heliocentric distances.

Oberon is the most distant major satellite. It spends part of its time outside of the uranian magnetosphere in the solar wind environment ([Cheng et al., 1991](#)). This difference between Oberon and the other uranian satellites offers a possible explanation for its H₂O ice bands following a different longitudinal pattern. Closer in, magnetospheric plasma may be the dominant remover of H₂O ice, chiefly from trailing hemispheres, while further out, micrometeoroid impacts may increase in relative importance, preferentially sputtering ice from leading hemispheres.

3.2. Carbon dioxide ice

Three narrow CO₂ ice absorption bands are seen in some of the spectra around 2 μ m. They are most apparent in Ariel spectra, where they were previously noted ([Grundy et al., 2003](#)). Here we report that these bands also appear, although less prominently, in spectra of Umbriel and Titania. [Fig. 3](#) shows the CO₂ ice absorptions, and a general trend of more CO₂ ice absorption on the satellites closer to Uranus (toward the bottom of the figure): no CO₂ absorption is evident in Oberon's spectrum, the strongest CO₂ absorptions are barely visible in Titania spectra, and the absorptions become progressively stronger in spectra of Umbriel and Ariel. Two additional CO₂ ice absorptions at 1.578 and 1.610 μ m also appear to be present in spectra of Ariel and possibly Umbriel and Titania. These absorptions are attributable to $2\nu_1 + 2\nu_2 + \nu_3$ and $\nu_1 + 4\nu_2 + \nu_3$ transitions in CO₂ ice, respectively ([Quirico and Schmitt, 1997](#); [Gerakines et al., 2005](#)). The $2\nu_1 + 2\nu_2 + \nu_3$ band looks like it is superimposed on the longer wavelength shoulder of a broader dip. That feature appears to be a stellar absorption that is slightly deeper in the Sun than in our solar analogues, resulting in incomplete cancellation. Additional, very weak, unidentified CO₂ ice absorptions from the laboratory spectra of [Hansen \(2005\)](#) are also visible in the regions enlarged in [Fig. 3](#). In general, SpeX does not quite spectrally resolve the cores of these CO₂ ice absorptions.

² NASA's first New Frontiers mission, bound for Pluto and the Kuiper belt ([Stern and Cheng, 2002](#); [Stern and Spencer, 2003](#)).

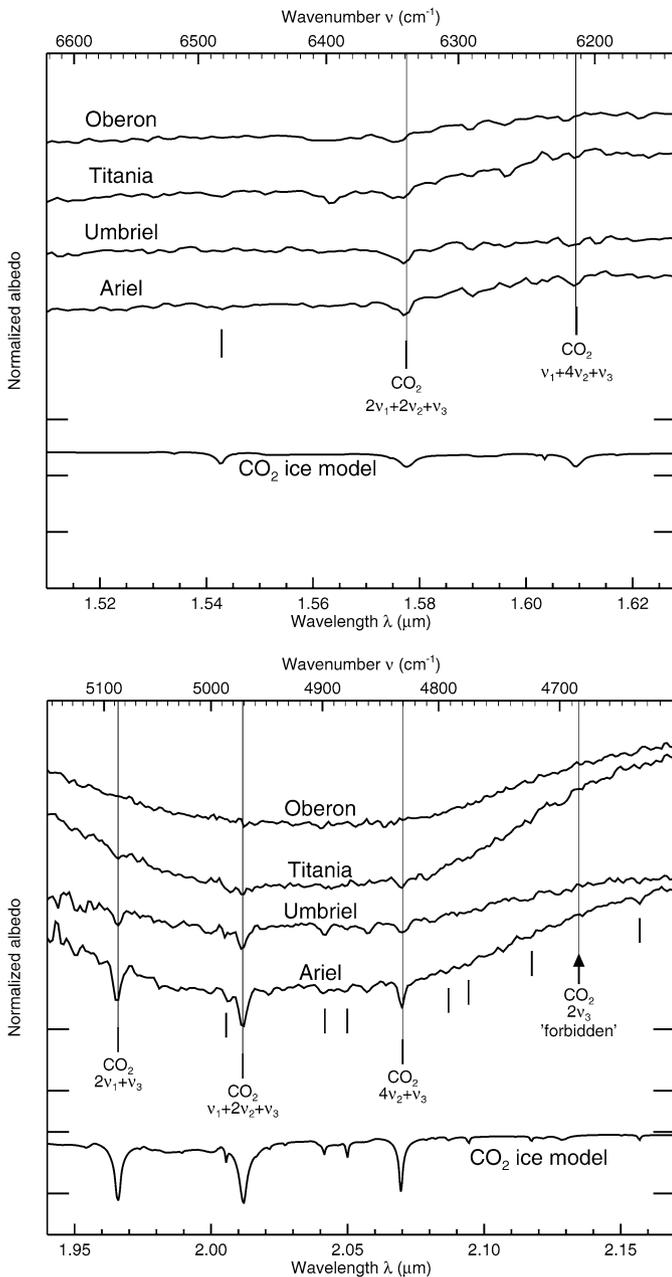


Fig. 3. Enlarged views of the 1.6 and 2 μm regions of the same satellite spectra as in Fig. 1, showing narrow CO_2 ice absorption features superimposed on the broad H_2O ice bands. Wavelengths of additional weak, unidentified CO_2 ice absorptions in the laboratory data of Hansen (1997, 2005) are indicated by vertical line segments. An arbitrarily scaled Hapke model spectrum (e.g., Hapke, 1993) based on those laboratory data is included at the bottom of each panel to give an idea of the sorts of band shapes expected for pure CO_2 ice, albeit at a warmer temperature (150 K). An arrow marks the location of the ‘forbidden’ $2\nu_3$ transition (Bernstein et al., 2005).

Our spectra show no apparent absorption features at 2.134 μm , where the ‘forbidden’ $2\nu_3$ overtone of CO_2 is reported by Bernstein et al. (2005) for CO_2 molecules mixed into H_2O ice (this behavior was also noted by B. Schmitt, personal communication, 2002). The absence of this band couples with the wavelengths and the narrow profiles of the other CO_2 bands to strongly favor pure CO_2 ice, in the sense of CO_2 molecules being associated with their own kind, as opposed to being iso-

lated in H_2O ice. That the CO_2 molecules are not isolated in H_2O ice is perhaps not surprising. There is some uncertainty about how long such a thermodynamically unstable situation would last. Laboratory evidence suggests that, at least at higher temperatures, $\text{H}_2\text{O}:\text{CO}_2$ ices mixed at the molecular level can reorganize themselves over time into segregated zones of pure phases, and thus lose the $2\nu_3$ absorption feature (Bernstein et al., 2005). It is unclear how rapidly such a reorganization could progress at the uranian satellites’ surface temperatures, but the apparent ability of H_2O to maintain its crystal structure against amorphization hints that such a phenomenon might not be out of the question.

The spatial distribution of CO_2 ice can place valuable constraints on its origin and evolution. Fig. 4 explores the longitudinal and planetocentric distribution of CO_2 ice on the four satellites by plotting the sums of the integrated areas of the three strongest CO_2 ice absorptions near 2 μm from each spectrum versus sub-observer longitude. The integrated areas for the three absorptions were computed using 1.957–1.962 μm and 1.969–1.974 μm for the continuum and 1.962–1.969 μm for the band interval of the $2\nu_1 + \nu_3$ CO_2 band. For the $\nu_1 + 2\nu_2 + \nu_3$ band we used 2.002–2.008 μm and 2.015–2.020 μm for the continuum and 2.008–2.015 μm for the band interval. For the $4\nu_2 + \nu_3$ band we used 2.062–2.068 μm and 2.072–2.078 μm for the continuum and 2.068–2.072 μm for the band interval.

These measurements show more CO_2 ice absorption on the trailing hemispheres of Ariel, Umbriel, and Titania, as well as declining CO_2 absorption with distance from Uranus. This distribution contrasts somewhat with reports from Galileo/NIMS observations of Europa, Ganymede, and Callisto (McCord et al., 1997, 1998; Hibbitts et al., 2000, 2003). Those observations of the 4.25 μm ν_3 fundamental band showed CO_2 to be present on both leading and trailing hemispheres, although more abundant at the center of the trailing hemisphere of Callisto. Also, the ν_3 band is shifted in wavelength from where it occurs in pure CO_2 ice, suggesting that in the jovian system CO_2 is trapped or bound up in some other material. The absorption was also observed to be geographically associated with darker materials (Hibbitts et al., 2000, 2003). Pure CO_2 ice in the jovian system would also not be thermally stable, adding weight to the interpretation that it is complexed with some other, less volatile material. The planetocentric trend was the opposite of what we see in the uranian system: the 4.25 μm band depth increases with distance from Jupiter (McCord et al., 1998). The ν_3 band of carbon dioxide has also recently been detected in saturnian satellite spectra by Cassini/VIMS (e.g., Buratti et al., 2005; Clark et al., 2005). As in the jovian system, the band appears to be shifted from the wavelength of the pure CO_2 ice transition, and may be associated with dark material there, as well.

It is important to realize that the 4.25 μm ν_3 fundamental absorption detected by NIMS and VIMS is intrinsically much stronger than the overtone and combination bands we are observing near 2 μm , by a factor of a thousand in absorption coefficients (e.g., Hansen, 1997, 2005). Accordingly, our detection of much the weaker bands requires considerably more condensed CO_2 than is indicated by the jovian and saturnian satellite observations. Simple radiative transfer models (e.g.,

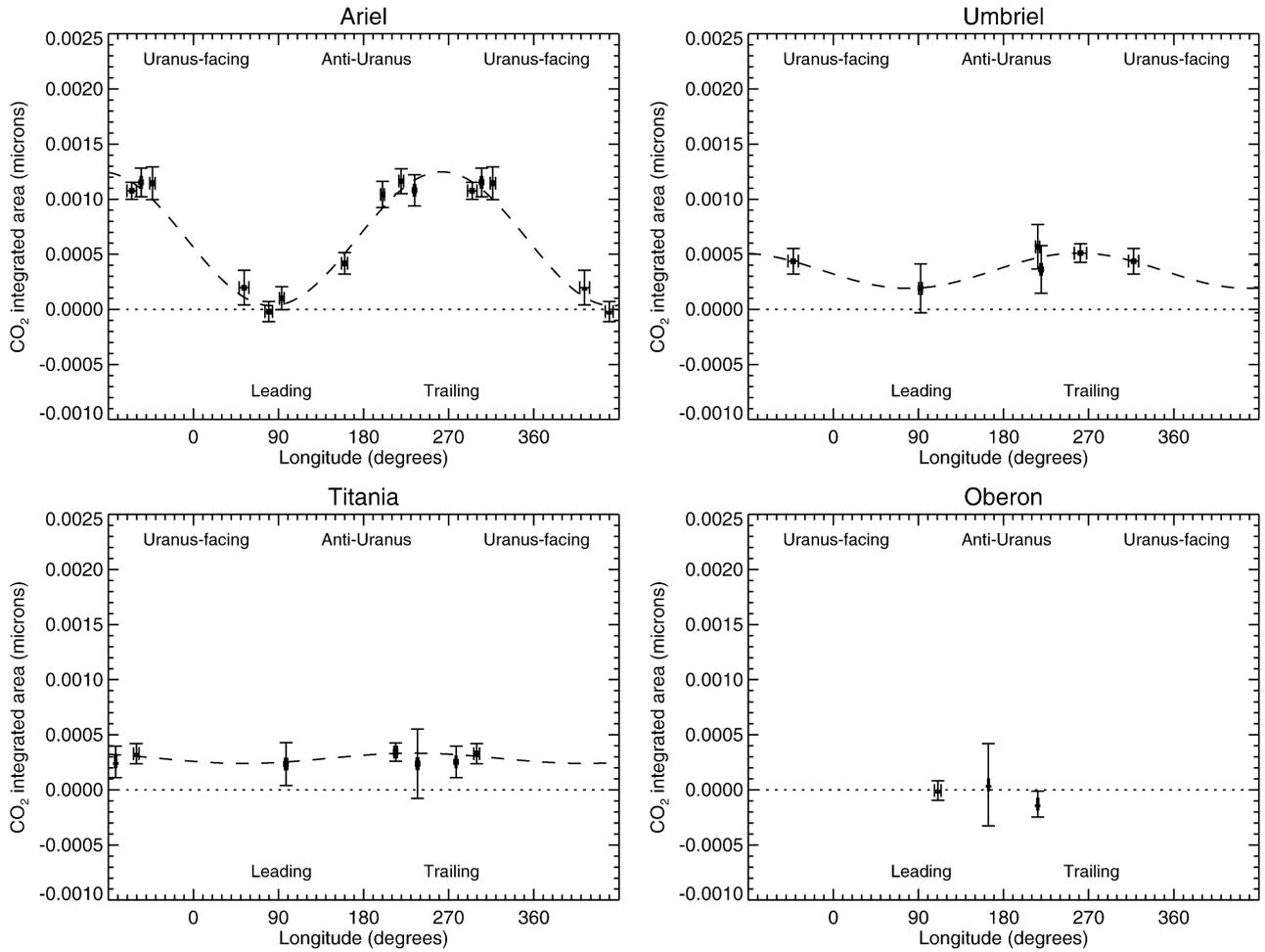


Fig. 4. Variation of total integrated area of the three strongest CO₂ ice bands near 2 μm . Sine fits are plotted as dashed curves. Both the integrated areas and their longitudinal variation decline with planetocentric distance. No CO₂ absorption was detected in the Oberon spectra.

Hapke, 1993; Grundy et al., 2003) can be used to get an idea of the quantities of CO₂ ice required by our observations. A CO₂ ice glaze uniformly covering Ariel's entire trailing hemisphere would have to be at least 5 μm thick to match the observed absorption bands. It would need to be thicker if the CO₂ ice were more geographically restricted. If the CO₂ ice were as segregated as possible from other materials, in a checkerboard-like pattern, at least 5% of the trailing hemisphere of Ariel would have to be composed of CO₂ ice.

In the following sub-sections, we consider possible sources and sinks of CO₂ on the satellites' surfaces, estimating rates where possible, and considering what longitudinal or planetocentric trends various processes might lead to in light of the observed CO₂ ice distribution. We begin with processes which can mobilize or eliminate CO₂ ice.

3.2.1. Processes which move or destroy CO₂

If the CO₂ ice on the satellite surfaces is not trapped in a less volatile material such as H₂O ice or a carbonaceous residue (as is suggested by the non-detection of the $2\nu_3$ band and by the absence of wavelength shifts relative to pure CO₂ ice measured in the laboratory), we can use its known vapor pressure curve (Brown and Ziegler, 1980) and sticking coefficient (Weida et

al., 1996) to consider its stability against sublimation over the course of a seasonal cycle in the uranian system. We used a thermophysical model (e.g., Spencer and Moore, 1992) to compute diurnal temperature histories as functions of latitude and season on the satellites, assuming plausible values of thermal inertia Γ , bolometric bond albedo A_B , and emissivity ϵ . From the computed temperature histories and the CO₂ ice vapor pressure curve, we computed sublimation histories for each latitude band. Integrating these histories over time, we arrived at seasonally averaged sublimation rates as a function of latitude. Inverting these rates gives the time to sublimate one gram of CO₂ ice per square centimeter. Examples are shown in Fig. 5.

The CO₂ sublimation rate is proportional to its vapor pressure, which is an extremely steep function of temperature at uranian satellite surface temperatures, so local seasonal sublimation rates tend to be dominated by the seasonal maximum temperatures achieved. Polar regions, which are bathed in decades of continuous sunlight each summer because of the extreme obliquity (98°), experience the highest peak temperatures and most rapid seasonal sublimation rates. This latitude dependence of seasonal sublimation rates will tend to drive accumulation of CO₂ at equatorial latitudes. But even at the satellites' equators, appreciable quantities of CO₂ ice could sublimate on

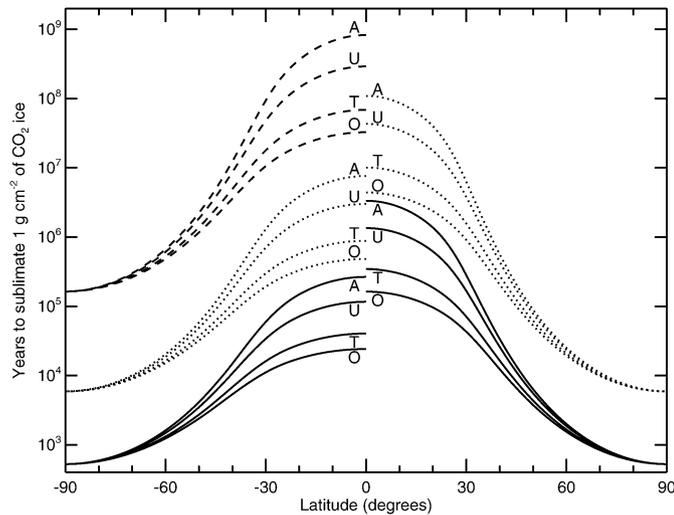


Fig. 5. Estimates of time in Earth years to sublimate one gram of CO_2 per cm^2 as a function of latitude for satellites as labeled by their initials. North and South latitudes mirror each other, so half of each curve is omitted to reduce clutter. On the left are models with $\Gamma = 10^4 \text{ erg cm}^{-2} \text{ s}^{-1/2} \text{ K}^{-1}$. Models with $\Gamma = 2 \times 10^4 \text{ erg cm}^{-2} \text{ s}^{-1/2} \text{ K}^{-1}$ are on the right. Solid curves are for $A_B = 0.5$, dotted curves for $A_B = 0.6$, and dashed curves for $A_B = 0.7$. Emissivity $\epsilon = 0.9$ was assumed for all models.

time scales much shorter than the age of the Solar System. Of the four satellites, Ariel spins fastest on its axis, with a 2.5 day period. Its faster rotation reduces the day–night temperature contrast, so for any set of thermophysical parameters, Ariel’s equatorial diurnal peak temperatures remain lower than on Umbriel, Titania, and Oberon, with their longer periods of 4.1, 8.7, and 13.5 days, respectively. As a result, CO_2 ice is most stable against sublimation at low latitudes on the faster rotating, inner satellites.

The extremely strong temperature dependence makes sublimation rates highly sensitive to assumed model parameters. For instance, increasing A_B from 0.5 to 0.6 reduces absorbed sunlight by 20%, leading to a few K drop in maximum temperatures. This small temperature drop drives an order of magnitude increase in the time scale to sublimate 1 g cm^{-2} of CO_2 ice, as shown in Fig. 5. Although we assumed plausible values for albedos, thermal inertias, emissivities, etc., in our computations of temperature histories, the extreme temperature sensitivity makes the sublimation numbers highly uncertain. Our models produce peak temperatures in the 70–80 K range. Peak temperatures from Voyager 2 thermal emission measurements were as high as 85 K (Hanel et al., 1986). If CO_2 ice temperatures get that high, sublimation will proceed much more rapidly than we have estimated in Fig. 5. The thermal emission observed by Voyager is dominated by the darkest regions which absorb the most sunlight as well as topographic lows that absorb thermal radiation from surrounding areas. Any CO_2 ice in these warmest regions would be mobilized very rapidly.

The net result of thermal mobilization will be that CO_2 ice migrates away from polar latitudes toward equatorial latitudes on relatively short time scales, of the order of hundreds to millions of years. In the equatorial regions, the ice will tend to accumulate in local cold traps: regions of higher albedo,

higher thermal inertia, higher emissivity, and/or higher elevation. There its texture will slowly evolve toward larger particles with rounded shapes, or even a compacted crust overlaying more porous ice below (e.g., Clark et al., 1983; Spencer, 1987; Eluszkiewicz et al., 1998; Grundy and Stansberry, 2000; Titus et al., 2001).

The above sublimation estimates do not account for the fate of CO_2 molecules after they are mobilized. Being unlikely to collide with other gas molecules, they will generally follow ballistic trajectories. Some will escape, while others will fall back to the surface and stick elsewhere. We estimate escaping fractions following Chamberlain and Hunten (1987, Eq. (7.1.5)). As with sublimation rates, the fraction escaping is extremely sensitive to temperature, and escape rates from the warmer poles greatly exceed escape rates from the cooler equatorial belts. For the thermal parameters described above, we find maximum polar escape fractions ranging from 1×10^{-9} (for Titania with $A_B = 0.7$) up to 7×10^{-4} (for Umbriel with $A_B = 0.5$). For equatorial latitudes, peak temperatures are lower and so maximum escape fractions are lower as well, from 1×10^{-10} to 3×10^{-4} for the same two cases, with $\Gamma = 10,000 \text{ erg cm}^{-2} \text{ s}^{-1/2} \text{ K}^{-1}$. For any particular set of thermal parameters, Umbriel has the highest escape fraction, followed by Ariel, because of their lower masses and gravitational accelerations. Since these two satellites show the most CO_2 ice, the escape of sublimated CO_2 to space does not seem to govern the planetocentric distribution of CO_2 . Even for the highest equatorial escape rates of any of our models, thermal escape of CO_2 would be extremely slow. For Umbriel with $\Gamma = 10,000 \text{ erg cm}^{-2} \text{ s}^{-1/2} \text{ K}^{-1}$ and $A_B = 0.5$, the time to sublimate 1 g cm^{-2} at the equator is about 10^5 years, of which $3 \times 10^{-4} \text{ g}$ escape to space, so the time to lose 1 g cm^{-2} to space is about 3×10^8 years. For other thermal model assumptions and for other satellites, the time scales are even longer, much longer than the age of the Solar System in most cases.

Charged particle sputtering can mobilize CO_2 , and will act at a much higher rate for CO_2 ice than it does for less volatile H_2O ice, assuming the two ices are segregated from one another (e.g., Johnson et al., 1983). Using Eq. (3.27) from Johnson (1990) with values from Table 3.4 in that book, we estimate that for typical uranian satellite surface temperatures and magnetospheric proton energies, the sputtering rate for CO_2 ice should be some 50–100 times higher than for H_2O ice. Cheng et al. (1991) estimate seasonal average sputtering rates for H_2O molecules from uranian satellite surfaces in their Table IV. We multiply their H_2O sputtering rates by 50–100 to get an idea of possible magnetospheric proton CO_2 sputtering rates for the uranian satellites, finding values in the $0.6\text{--}1.4 \times 10^{-8} \text{ g cm}^{-2} \text{ year}^{-1}$ range for Ariel and Umbriel, dropping to $2 \text{ to } 4 \times 10^{-10} \text{ g cm}^{-2} \text{ year}^{-1}$ for Titania and as low as $0.5\text{--}1 \times 10^{-10} \text{ g cm}^{-2} \text{ year}^{-1}$ for Oberon. As before, these values can be inverted to get an idea of how long it would take for one gram of CO_2 ice to be sputtered from a square centimeter. For Ariel and Umbriel, these time scales are in the 100 Myr range, increasing to more than a billion years for Titania and Oberon.

Sputtered molecules are ejected with more kinetic energy than thermally mobilized molecules generally have, so they can more readily escape a satellite's gravitational well. The energy spectrum of sputtered CO₂ molecules is uncertain, but crude estimates can be made by assuming it is similar to the energy distribution of sputtered H₂O molecules, which goes approximately as $U'(E + U')^{-2}$, where U' is a measure of binding energy and E is the kinetic energy of ejection (e.g., Johnson, 1998). The appropriate value of U' for CO₂ ice is unknown. Assuming an H₂O-like $U' = 0.05$ eV (Haring et al., 1984) the fraction of molecules sputtered from the surfaces of the satellites at greater than escape velocity would be 0.85, 0.40, 0.44, 0.26, and 0.29 from Miranda through Oberon. If U' is as low as the 0.009 value reported by Haring et al. (1984) for CO ice, the escape fractions would be 0.51, 0.11, 0.13, 0.06, and 0.07. These values are generally smaller than the Cheng et al. (1991) estimated escape fractions for H₂O molecules, since CO₂ molecules are more massive (44 AMU as opposed to 18 AMU for H₂O). Combining the escape fractions with sputtering rates as estimated in the previous paragraph, we find time scales for one gram of CO₂ ice to be sputtered and lost to space to be in the few hundred million year range for Ariel and Umbriel, but longer than the age of the Solar System for Titania and Oberon.

Sputtering will gradually remove CO₂ ice from the low latitude cold traps where thermal mobilization drives it, at least on the trailing hemispheres of the inner satellites, where magnetospheric sputtering rates are highest and the molecules are less tightly bound, gravitationally, because the satellites are less massive (see Table 1). The fact that CO₂ ice is observed to be most abundant in exactly these regions where it should be removed most rapidly suggests sputtering by magnetospheric protons is not controlling the distribution of CO₂ ice on the satellites' surfaces, and that whatever is controlling that distribution acts on shorter time scales.

Sputtering by solar UV photons could also mobilize CO₂ on satellite surfaces (e.g., Johnson, 1990). UV sputtering should affect all longitudes and all uranian satellites equivalently, so although it may be important for increasing mobility of CO₂ molecules, it cannot explain the observed spatial distribution of CO₂ ice. Another source of energy which can lead to sputtering is micrometeorite impacts, which, if they come from outside of the Uranus system, preferentially strike the satellites' leading hemispheres (Zahnle et al., 2003). The abundance of dust-sized impactors at 20 AU is not well known, but could be sufficient to play a role in mobilizing CO₂ ice from the leading hemispheres (e.g., Eviatar and Richardson, 1986).

Finally, CO₂ molecules can be destroyed by energetic UV or charged particle radiation. Gerakines and Moore (2001) estimated G values of 1.1 and 8.1 CO₂ molecules destroyed per 100 eV incident as protons or UV photons, respectively, via the reaction path energy + CO₂ → CO + O. A UV photon with wavelength shorter than 2275 Å can dissociate a CO₂ molecule (Gerakines et al., 1996; Delitsky and Lane, 1998). From SORCE spectra (e.g., Rottman et al., 2004), we estimate about 0.07% of the solar flux to be at or below that wavelength. Accounting for the geometry of the Uranus system, we obtain seasonal average UV fluxes ranging from 0.5 through

0.8 erg cm⁻² s⁻¹ for equatorial through polar latitudes on the uranian satellites. These fluxes correspond to destruction of the order of 10¹⁸ CO₂ molecules, or ~10⁻⁴ g cm⁻² yr⁻¹, if all of the UV photons are absorbed within the CO₂ ice. This is a very high destruction rate, compared with other processes we have considered.

Magnetospheric charged particles can also dissociate CO₂ molecules. Using the Gerakines and Moore (2001) G value of 1.1 molecule per 100 eV, along with dose rates from Cheng et al. (1991, Fig. 23), we estimate radiolytic destruction rates in the 10⁻¹⁰–10⁻⁷ g cm⁻² yr⁻¹ range for Oberon through Miranda. These rates are much lower than the photolytic destruction rates estimated in the previous paragraph, but are comparable to the magnetospheric sputtering rates estimated previously.

Of course, not all CO and O resulting from photolysis and radiolysis of CO₂ will be lost. Some will recombine to make new CO₂ molecules while others will react with water or other locally available materials to form new molecules, some of which could be quite complex (e.g., Benit et al., 1988; Delitsky and Lane, 1997, 1998). Indeed, it seems likely that a carbon cycle much like the one envisioned by Johnson et al. (2005) for the jovian system also operates on the surfaces of the uranian satellites, with carbon atoms repeatedly recycled between CO₂ and other molecules such as carbonic acid (H₂CO₃) and formaldehyde (H₂CO).

The relatively high destruction and loss rates estimated in this section, especially from UV photolysis as well as from sputtering and radiolysis, suggest that there must be a recent, or more likely an ongoing source for the observed CO₂ ice. Accordingly, we next turn our attention to possible sources of CO₂.

3.2.2. Possible sources of CO₂

Out-gassing from satellite interiors is a possible source for CO₂. Ariel, with its relatively deep CO₂ absorption bands, also shows clear evidence of past geologic activity (Plescia, 1987), but Umbriel, with the next strongest CO₂ bands, shows no evidence for recent activity, at least not on the hemisphere observed by Voyager (Plescia, 1989). Titania, with still less observed CO₂, may have some relatively youthful provinces, but its surface is mostly ancient (Zahnle et al., 2003). Although Voyager was unable to view more than about half of each satellite's surface, the presence of CO₂ ice does not seem to correlate particularly well with the existence of less-cratered, younger regions or with evidence of recent geological activity. Either could be an indicator of ongoing or recent release of volatiles from the satellites' interiors. An out-gassing source for the observed CO₂ ice thus does not look especially promising.

Impactors could deliver CO₂ to the uranian satellites, or could release it from their sub-surfaces. If exogenic delivery or impact exhumation were dominated by frequent, smaller impactors, they would tend to distribute CO₂ through the Uranus system according to a consistent pattern, and the balance between delivery and loss rates would govern the observed planetocentric and longitudinal distributions. Infrequent, larger impactors could produce a more random distribution pattern, in which a local high abundance of CO₂ might simply point to

a single recent impact event. Zahnle et al. (2003) estimated cometary cratering rates on the uranian satellites. Owing to gravitational focusing, cratering rates per unit area increase with proximity to Uranus, being something like a factor of 4 or 5 higher on Ariel than on Oberon, and impactors strike with average speeds a factor of almost 2 higher on Ariel than on Oberon (Zahnle et al., 2003). Impactors from outside the Uranus system would preferentially strike the satellites' leading hemispheres, so the observation that CO₂ ice is predominantly on the trailing hemispheres would require mobilization and redistribution, followed by preferential removal from the leading hemispheres, possibly by dust impact driven sputtering.

CO₂ molecules could form through radiolytic action on the satellites' surfaces, with oxygen coming from the ubiquitous H₂O ice, and carbon arriving via implantation of magnetospheric ions (Strazzulla et al., 2003). A magnetospheric source of carbon ions would neatly explain the observed distribution of CO₂ ice. However, the Voyager 2 spacecraft did not detect carbon ions in the magnetosphere during its Uranus flyby in 1986 (Bridge et al., 1986; Krimigis et al., 1986), making a magnetospheric source of carbon ions seems improbable. However, magnetospheric delivery of larger carbonaceous particles from the uranian rings remains a possibility, as mentioned previously.

There is also no shortage of dark and presumably carbonaceous material already on the satellites' surfaces, as evidenced by their low visible albedos. Recent investigations of radiolysis at interfaces between carbonaceous materials and H₂O ice (Mennella et al., 2004; Gomis and Strazzulla, 2005) show that CO₂ can be produced from purely local materials, even when the carbon and oxygen are not initially present in a single phase. The low albedos of the satellites' surfaces, along with temperatures too cold for thermal segregation of H₂O ice as well as abundant craters indicative of impact gardening, suggest relatively intimate mixing between H₂O ice and dark particles. In one experiment by Gomis and Strazzulla (2005), H₂O ice deposited on a carbonaceous substrate was irradiated at 80 K with 200 keV Ar⁺ ions. CO₂ was efficiently produced at the interface. It would be interesting to see if proton, electron, or even UV irradiation produces similar results. Any energy source capable of breaking bonds in H₂O ice or in the dark carbonaceous materials, and mobilizing radicals containing O or C should be able to contribute to this process.

Whether or not radiolysis can produce CO₂ rapidly enough to replace CO₂ lost to UV photolysis is an important question to consider. Prior to saturation in the Gomis and Strazzulla (2005) experiment, each 200 keV Ar⁺ ion produced about 30 CO₂ molecules, or 7 keV per CO₂ molecule produced. This is orders of magnitude more energy than the ~12 eV per CO₂ molecule destroyed according to the photolytic *G* value from Gerakines and Moore (2001). From these numbers, radiolytic production rates would be far below photolytic destruction rates. For production not to be overwhelmed by photolytic destruction, much more efficient production pathways are necessary. It is possible that lower energy protons or electrons are more effective at producing CO₂ from the available precursors, and it is also possible that the expected granular texture of the ice

and carbonaceous material enhances production of CO₂, because of its large surface area. Additionally, photolysis may not be a purely destructive process. We are not aware of laboratory studies of photolytic production of CO₂ from H₂O ice and carbonaceous precursors at uranian satellite surface temperatures, but it probably can occur (e.g., Bernstein et al., 1995; Wu et al., 2002). Photolytic production would not, however, be able to explain the observed longitudinal distribution of CO₂ ice. Another possibility is that more of the satellites' surfaces are involved in production than in destruction. Such a configuration is plausible, since CO₂ will tend to accumulate in localized cold traps, owing to its high thermal mobility, unlike the global distribution of the non-volatile precursors: H₂O ice and carbonaceous solids. As noted previously, the spectra are consistent with CO₂ being confined to a relatively small fraction of the surface area, as little as 5%.

3.2.3. Discussion

The planetocentric and longitudinal distributions of CO₂ and H₂O ices in the uranian system offer valuable clues regarding their origin and fate. Having considered candidate sources and sinks on the satellite surfaces, we now review the trends expected from these mechanisms and compare them with the observational data, which shows CO₂ ice to be most abundant on the inner satellites, and especially on their trailing hemispheres, and H₂O ice to be more abundant on leading hemispheres from Ariel through Titania, but on the trailing hemisphere of Oberon.

Sublimation mobilizes CO₂ relatively rapidly at seasonal maximum uranian satellite surface temperatures, which are higher at polar latitudes, so CO₂ accumulates at low latitudes. Relatively little is lost to space. At low latitudes, sublimation rates are much lower, but still rapid enough to drive CO₂ ice to accumulate in local cold traps, and to evolve in texture. Sublimation does not depend on longitude, and has no real effect on the much less volatile H₂O ice. Voyager images of Umbriel show a bright equatorial feature ('Wunda') near the center of the trailing hemisphere. It is tempting to speculate that CO₂ ice is concentrated there, however, similar low-latitude, bright features are not evident on the trailing hemispheres of Ariel or Titania.

Impactors preferentially strike leading hemispheres and inner satellites. Impacts could dredge up fresh H₂O ice from the subsurfaces, primarily of the leading hemispheres of the inner satellites, consistent with the observations of deeper H₂O ice bands there. An impact source of CO₂ could be consistent with the observed planetocentric trend, but not with the observed longitudinal trends. An impact loss mechanism could explain the observed CO₂ longitudinal trends, but not the planetocentric trend. Impacts apparently cannot explain the observed CO₂ ice distribution, at least not without the help of additional mechanisms.

Magnetospheric charged particle sputtering preferentially mobilizes CO₂ and H₂O ices from the trailing hemispheres of the inner satellites, exactly those regions where CO₂ ice is observed. Sputtering rates are estimated to be well below sublimation rates for CO₂, but are probably the dominant sink for H₂O ice. A much greater fraction of sputtered CO₂ mole-

cules should escape to space, compared with sublimated CO₂ molecules. Correlations with indicators of recent geological activity in Voyager images are not evident, arguing against an out-gassing source of CO₂ from the satellite interiors.

Radiolytic production of CO₂ driven by magnetospheric irradiation should act most rapidly on the trailing hemispheres of the inner satellites, because magnetospheric densities are higher closer to Uranus. The longitudinal and planetocentric pattern of radiolytic production is consistent with the observed distribution of CO₂, and appears to be the best candidate source for CO₂ on the satellites' surfaces. The fact that our estimated rates of radiolytic production are considerably below the estimated photolytic destruction rate is troubling. This discrepancy may point to problems in our estimated production rates or could imply that much more surface area is involved in production than is actually covered with CO₂ ice, in order to maintain an equilibrium. The idea that CO₂ ice may only cover a small fraction of the surface area of the satellites' trailing hemispheres is at least consistent with the spectral evidence.

Finally, it seems probable that CO₂ ice on the uranian satellites' surfaces participates in a radiolytic and photolytic carbon cycle (e.g., Johnson et al., 2005). If so, carbonic acid is also likely to be present, and might be detectable by means of its characteristic infrared absorption bands. With hydrogen being more readily lost to space than heavier carbon and oxygen, chemistry at the satellites' surfaces probably tends to be oxidizing and acidic. Such a surface environment may help explain why NH₃, an alkali, has not been convincingly detected, despite its presence being suggested by geomorphologic evidence (Croft and Soderblom, 1991; Kargel et al., 1991). A tentative report of NH₃ in a spectrum of Miranda (Bauer et al., 2002) has yet to be confirmed. If NH₃ were confirmed to be present on the surface of a uranian satellite, the chemical implications would be quite remarkable.

4. Conclusion

Seventeen nights of IRTF/SpEx observations of the 4 largest uranian satellites reveal the existence of CO₂ ice on Umbriel and Titania for the first time, as well as confirming its presence on Ariel. The observations also provide information about the longitudinal and planetocentric distributions of CO₂ and H₂O ices. These distributions likely result from equilibria between source and sink mechanisms, and can tell us something about the relative importance of the processes involved. Magnetospheric sputtering removal of H₂O ice from trailing hemispheres of Ariel, Umbriel, and Titania is consistent with the observation of deeper H₂O ice absorption bands on the leading hemispheres of those satellites, with diminishing leading-trailing asymmetries with planetocentric distance. For Oberon, magnetospheric sputtering is greatly reduced, and the leading-trailing asymmetry appears to be reversed. The deeper H₂O ice bands on Oberon's trailing hemisphere may point to an effect of micrometeoroid sputtering that only becomes important where magnetospheric effects become negligible. Radiolytic production of CO₂, driven by magnetospheric charged particle bombardment, appears most consistent with the observation of

deeper CO₂ ice absorption bands on the trailing hemispheres of Ariel, Umbriel, and Titania, as well as declining CO₂ absorption with planetocentric distance and its absence from the surface of Oberon.

Thermally-driven processes such as grain growth, volatile transport, sintering, and solar gardening can be expected to influence the texture and distribution of CO₂ ice on the surfaces of the uranian satellites. The seasonally-averaged distribution of sunshine should drive CO₂ ice away from polar latitudes. Low latitude regions which radiatively cool fastest will cold-trap CO₂ ice, and if enough ice accumulates to raise the albedos of these regions, the process will be self-reinforcing. Uppermost surfaces of CO₂ ice particles can radiatively cool faster than subsurface depths where solar energy is absorbed, so sublimation will tend to preferentially remove CO₂ from below the uppermost surface. Loss to space driven primarily by sputtering by UV photons and magnetospheric charged particles limits the residency time of CO₂ ice in these cold traps. Additional limits on CO₂ accumulation come from UV photolysis, which is expected to break CO₂ bonds at a relatively high rate, leading to a chemical cycle in which carbon, oxygen, and hydrogen are repeatedly combined into new species and broken apart again, forming a complex mixture which should include CO₂ as well as species like carbonic acid and formaldehyde. More detailed comparisons between leading hemispheres (which see relatively little magnetospheric radiolysis) and trailing hemispheres should be able to shed more light on the relative importance of radiolysis and photolysis.

If radiolysis by magnetospheric particles is indeed the source of CO₂ ice on the uranian satellites, production rates on Miranda should be even higher than on Ariel. However, Miranda's lower gravity will allow a much larger fraction of sublimated and sputtered CO₂ molecules to escape. Spectral observations of Miranda could provide useful constraints on the balance between sublimation and radiolytic production.

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